Far-infrared constraints on the contamination by dust-obscured galaxies of high-z dropout searches

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ABSTRACT

The spectral energy distributions (SED) of dusty galaxies at intermediate redshift may look similar to very high-redshift galaxies in the optical/near infrared (NIR) domain. This can lead to the contamination of high-redshift galaxy searches based on broad-band optical/NIR photometry by lower redshift dusty galaxies because both kind of galaxies cannot be distinguished. The contamination rate could be as high as 50%. This work shows how the far-infrared (FIR) domain can help to recognize likely low-z interlopers in an optical/NIR search for high-z galaxies. We analyze the FIR SEDs of two galaxies that are proposed to be very high-redshift (\(z > 7\)) dropout candidates based on deep Hawk-I/VLT observations. The FIR SEDs are sampled with PACS/Herschel at 100 and 160 \(\mu\)m, with SPIRE/Herschel at 250, 350 and 500 \(\mu\)m and with LABOCA/APEX at 870 \(\mu\)m. We find that redshifts \(z > 7\) would imply extremely intense FIR SEDs (with dust temperatures >100 K and FIR luminosities >10\(^{13}\) \(L_\odot\)). At \(z \approx 2\), Instead, the SEDs of both sources would be compatible with typical ultra luminous infrared galaxies or submillimeter galaxies. Considering all available data for these sources from visible to FIR we re-estimate the redshifts and find \(z \approx 1.6–2.5\). Owing to the strong spectral breaks observed in these galaxies, standard templates from the literature fail to reproduce the visible-to-near-IR part of the SEDs even when additional extinction is included. These sources strongly resemble dust-obscured galaxies selected in Spitzer observations with extreme visible-to-FIR colors, and the galaxy GN10 at \(z = 4\). Galaxies with similar SEDs could contaminate other high-redshift surveys.

Key words. galaxies: distances and redshifts – dust, extinction – gravitational lensing: weak – galaxies: high-redshift

1. Introduction

Observing galaxies up to very high-redshifts allows us to study directly the formation and evolution of structures in the expanding Universe. Finding galaxies at ever higher redshifts has therefore become one of the main areas of extragalactic astronomy. The most common technique is to use known broad features in the spectral energy distributions (SEDs) of galaxies to identify high-redshift sources in deep optical and near-infrared (NIR) multi-band observations. In particular the Lyman break is widely used to select sources by redshift, noting their disappearance in bands below a given wavelength, the so-called dropout technique (Steidel et al. 1996). With this technique and state-of-the-art telescopes and instruments it is now possible to select sources that are good candidates for being at the end or within the epoch of reionization (Richard et al. 2006; Zheng et al. 2009; McLure et al. 2010; Wilkins et al. 2010; Oesch et al. 2010; Bouwens et al. 2010b,a).

Low-redshift galaxies, however, can have very steep SEDs that resemble a break in the UV/optical/NIR. This can lead to contamination of the dropout selection of very high-z galaxies, and consequently to erroneous estimates of the star-formation rate density, stellar masses, and others, although these effects...
are currently difficult to quantify. Objects such as these have been found and discussed by several authors (see e.g. Dickinson et al. 2000; Mobasher et al. 2005; Schaerer et al. 2007; Dunlop et al. 2007; Chary et al. 2007; Capak et al. 2011). Confirming the photometric redshifts of high-z galaxies by identifying spectral lines is challenging because the sources are generally too faint for spectroscopic follow-up observations or because they may intrinsically lack Lyα emission (but see Vanzella et al. 2011).

The recent developments of space far-infrared (FIR) instrumentation offer new perspectives in this domain. In particular, with the advent of the Herschel Space Observatory it is now possible to sample the FIR part of the SEDs, where the thermal dust emission dominates. The shape of the FIR SED universally looks like a broad bump and can be used to further constrain the optical/NIR photometric redshifts. Although the wavelength of the FIR SED peak also depends on the dust temperature, the limited range of average temperatures observed so far in galaxies (between 20 and 60 K averaged over the entire galaxies, see, e.g., Kovács et al. 2006; Magnelli et al. 2010; Magdis et al. 2010; Wardlow et al. 2011) can be used as a prior and makes it possible to distinguish between intermediate (z < 3) and very high-redshifts (z > 6).

Far-infrared observations of high-z candidates are also essential to characterize their star forming and dust properties and thus interpret correctly their contribution to the cosmic history of star-formation and reionization.

Recently Laporte et al. (2011) identified ten z > 7 candidates in the field of the cluster Abell 2667 using photometric dropout criteria based on deep observations with HAWK-I on the ESO Very Large Telescope (VLT). Comparing their results to other studies and in particular to the WIRCAM Ultra Deep Field Survey (WUDS; Pello et al. in prep.), which is based on deeper optical observations blueward of the i-band, they estimated that 50–75% of these candidates could in fact be lower redshift interlopers. Here, we study two galaxies of this sample that are clearly detected by Herschel, namely the sources named “z1” and “Y5”. Our goal is to determine whether they could be interlopers and to understand their nature. The redshift probability distributions of these two sources derived by Laporte et al. (2011) from SED fitting to deep optical/NIR photometry show a prominent peak at z = 7.6 and 8.6 respectively. However, a secondary peak at lower redshift around z ~ 2 indicates that they could be interlopers as well. Laporte et al. (2011) also noted that the 24 μm detection of z1 with MIPS/Spitzer (Y5 is outside the Spitzer map) seems difficult to reconcile with the high-z solution. We use new Herschel and LABOCA observations of Abell 2667 to reconstruct the FIR part of their SEDs. We can thus further constrain their redshifts and study their physical properties.

The layout of the article is as follows: Sect. 2 gives a presentation of the observations and data analysis. In Sect. 3 the FIR part of the SEDs is analyzed. In Sect. 4 the complete SEDs are used to estimate the redshifts and discuss the physical properties of the two galaxies. In Sect. 5 we compare the two galaxies to other similar galaxies found in the literature. Section 6 gives the conclusions. We assume a Λ-cosmology with \( H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}, \Omega_M = 0.3 \) and \( \Omega_\Lambda = 0.7 \).

2. Observations and data analysis

2.1. Observations and reduction

Herschel observations were obtained in the framework of the Herschel Lensing Survey (HLS) described by Egami et al. (2010). They include PACS data at 100 and 160 μm, and SPIRE observations at 250, 350, and 500 μm. The data reduction was made with the HIPE software as described by Rex et al. (2010) and Rawle et al. (2010).

The large APEX Bolometer Camera (LABOCA Siringo et al. 2009) is a bolometer array operating at 870 μm that is mounted on the APEX telescope in the desert of Atacama, Chile (Güsten et al. 2006). The LABOCA observations were conducted during the summer 2010. The cluster was mapped in spurious mode during 30 h, covering a circular field of ~6° in radius. The data were reduced with the BoA1 software. The noise is not uniform over the map and the RMS is in the range 1.1–3.0 mJy, the highest values are reached at the edges of the map.

We also obtained a VLA 1.4 GHz continuum map of Abell 2667 (PI: Ivison) with an RMS of 46 μJy.

2.2. Analysis

The astrometry of all maps was corrected to align them with the VLT Ks image. All optical dropout sources of Laporte et al. (2011) were inspected in the Herschel and LABOCA images. Two of them, z1 and Y5, are detected in several FIR bands. Because IRAC/Spitzer and MIPS/Spitzer data are available for z1, the source can be followed from one band to the next one by increasing wavelength despite the decreasing resolution. Its identification is therefore robust.

For Y5 there is a larger gap in the SED owing to the lack of data between 8 μm and 100 μm, and due to the fact that it lies at the noisy edges of the 100 and 160 μm maps, where it is not detected. However, Y5 is the only source detected at 4.5 μm within a radius of 3′′ (i.e., ~1/3rd of the 250 μm beam radius) around the 250 μm peak, its identification with the SPIRE detection is therefore very likely.

The fluxes are measured at the positions of the two galaxies by PSF fitting in apertures with a radius equal to FWHM/3, where FWHM is the PSF full width at half maximum, i.e., 5.6′′, 11.3′′, 18.1′′, 24.9′′, 36.6′′ and 22.5′′ from 100 to 870 μm. The last (LABOCA) FWHM corresponds to the APEX beam convolved by a Gaussian of 12′′. The sources were deblended from the neighboring sources by subtracting PSFs at the positions of the neighbors derived from the 250 μm map. Observations at these wavelengths with these resolutions are affected by source confusion. As a consequence a measured flux cannot be directly interpreted as the true flux of a single underlying source. A correct treatment of the effect of source confusion on flux measurements (a.k.a. flux “deboosting”) requires a prior knowledge of the source counts toward low fluxes at the given wavelength. We followed the method presented by Crawford et al. (2010) based on a Bayesian analysis. For the prior source counts we extrapolated toward low fluxes the results of Berta et al. (2010) for PACS bands, Oliver et al. (2010) for SPIRE bands and Coppin et al. (2006) for the LABOCA band.

Blending affects z1 photometry at λ ≥ 250 μm and Y5 photometry at λ ≥ 500 μm. And the effect of deboosting is small (<20%), except for the 870 μm measurement of Y5, which corresponds to a 2.6σ signal and which we chose to consider as a tentative detection. The deboosted flux of Y5 at 870 μm is 1.8 ± 1 mJy for a measured flux of 2.5 ± 0.95 mJy.

None of the two sources are detected in the VLA map. The measured FIR fluxes of the sources as well as their optical-to-near-IR photometry from Laporte et al. (2011) are listed in Table 1. Thumbnails of the Herschel and LABOCA bands centered at the source positions as well as FIR SED fits are

\[ \frac{1}{http://www.apex-telescope.org/bolometer/laboca/boa/} \]
shown in the Fig. 1. Y5 is close to the border of the PACS maps where the noise is higher, hence the high upper limits.

### 3. Analysis of the FIR SEDs

The following models or templates were fitted to the FIR measurements (cf. Fig. 1):

- a modified black-body SED parameterized as described by Blain et al. (2003), with emissivity fixed to $\beta = 1.5$ and the Wien correction parameter $\nu = 2.9$. These values are adapted to submillimeter galaxies (SMGs; Chapman et al. 2005) and local ultraluminous infrared galaxies (ULIRGs; Dunne et al. 2000; Blain et al. 2003). The free parameters are the total FIR luminosity, $L_{\text{FIR}}$, defined as the luminosity emitted in the range $8-1000 \mu m$, and the dust temperature, $T_d$;
- the 105 galaxy templates built by Chary & Elbaz (2001). The templates are fitted without rescaling$^2$;
- ULIRG templates built by Vega et al. (2008), with a scaling parameter, $L_{\text{FIR}}$;
- the starburst, Seyfert, and active galactic nuclei (AGN) templates of Polletta et al. (2007), with a scaling parameter, $L_{\text{FIR}}$;
- templates built by Michałowski et al. (2010a,b) to fit high-redshift galaxies with detected but poorly sampled submm emission, with a scaling parameter, $L_{\text{FIR}}$;
- the SED fit to the observations of SMMJ2135-0102 (Swinbank et al. 2010; Ivison et al. 2010) with a scaling parameter, $L_{\text{FIR}}$.

The fit was performed by finding the maximum likelihood assuming Gaussian probability distributions for the measurements. When there is no detection, the $3\sigma$ value is used as a hard upper limit, i.e., the probability is assumed to be uniform in the [0, $3\sigma$] interval and zero outside. The redshifts are fitted to the solutions derived by Laporte et al. (2011) from the optical/NIR photometry, i.e., $z = 1.8$ and 7.6 for $z_1$ and $z = 1.7$ and 8.6 for Y5. The corresponding magnification factors are $\mu = 1.12$ and 1.17 for $z_1$ and $\mu = 1.04$ and 1.15 for Y5. The MIPS/Spitzer 24 $\mu m$ flux of $z_1$ (Y5 has no 24 $\mu m$ data available) was taken into account to fit various galaxy templates, but ignored to fit the modified black-body because it is most likely dominated by polycyclic aromatic hydrocarbons (PAHs).

For both sources we find reasonable fits at low-redshift for the modified black-body and the various galaxy templates. A ULIRG template from the Vega et al. (2008) library, a submillimeter-detected galaxy template from the Michałowski et al. (2010a,b) library and a Seyfert template from the (Polletta et al. 2007) library are able to reproduce the 24 $\mu m$ emission of $z_1$. The modified black-body model gives dust temperatures of $34$ and $36 K$, for $z_1$ and Y5 respectively, which are typical values for integrated dust temperatures in LIRGs. The infrared luminosities, $L_{\text{IR}}$, are in the range $(1.2-1.7) \times 10^{12} L_\odot$ and $(3.1-4.7) \times 10^{12} L_\odot$ for $z_1$ and Y5 respectively. There is a noticeable agreement in $L_{\text{IR}}$ between the modified black-body model and the various templates. These galaxies would therefore be typical ULIRGs/SMGs at $z \sim 2$. This is consistent with the general picture of galaxy evolution now widely observed, i.e., that the contribution of ULIRGs to the cosmic SFR is expected to peak at $z \sim 2$ where it should be comparable to that of the more “normal” galaxies (see e.g. Murphy et al. 2011).

For the high-redshift solutions ($z > 7.5$), instead, the modified black-body requires for both sources very high dust temperatures, i.e., $105$ and $129 K$ for $z_1$ and Y5, respectively. Such high temperatures averaged over an entire galaxy are extreme. This can be seen from the impossibility to find any good fit in the different template libraries, which were built from observed galaxies. However, the dust properties of galaxies at such high-redshifts are unknown and dust temperatures above $100 K$ cannot be ruled out. This would imply that the FIR luminosities are on the order of $0.5$ and $1.5 \times 10^{14} L_\odot$ for $z_1$ and Y5, respectively, i.e., both sources would be classified as hyper luminous infrared galaxies (HyLIRGs). While the nature of HyLIRGs is still a matter of debate (e.g. Ruiz et al. 2010) and their density at very high-redshift is not well known, they are extreme sources with a lower number density than ULIRGs.

The radio continuum upper limits are too high to constrain the SED fitting. We find that for Y5 only at the low-redshift solution ($z \sim 2$) the Chary & Elbaz template is close to the $3\sigma$ upper limit.

Thus, in summary, by comparing the FIR photometry to known galaxy SEDs and by taking into account the expected temperature and luminosity range of high-redshift galaxies, the very high-redshift solutions derived from the optical/NIR photometry seem to be less likely than the low-redshift solution. The two sources are most likely typical ULIRGs at $z \sim 2$. This result puts strong constraints on the optical/NIR analysis, which gave a much higher probability to the very high-redshift solution when no prior luminosity function was taken into account. The FIR data alone, however, cannot be used to derive any accurate photometric redshift because of the redshift-temperature degeneracy.

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2 We found that this additional scaling parameter was not required to obtain good fits.

3 To our knowledge there was only one such extreme case reported so far: the host of the lensed quasar APM08279 at $z = 3.9$ (Weiß et al. 2007; Riechers et al. 2009). It requires two dust components at 75 and 220 K and the highest temperatures most likely result from heating by the powerful quasar.
4. Analysis of the complete SED from visible to FIR

We will now examine all data from the visible to the FIR ranges to improve the redshift estimate of our galaxies and to examine the nature and physical properties of these sources.

4.1. Method

To model the SED of the two sources we used a modified version of the Hyperz photometric redshift code of Bolzonella et al. (2000) described in Schaerer & de Barros (2009). Non-detections are treated as the usual case 1 of Hyperz, i.e., the flux in these filters is set to zero, with an error bar corresponding to the flux at 1σ level. The basic spectral templates are taken from the Bruzual & Charlot models (Bruzual & Charlot 2003), computed for a variety of star-formation histories and metallicities. Although applicable only to a limited part of the spectrum, we use these templates here to constrain redshift, extinction, and stellar mass in particular. For the Bruzual & Charlot templates we consider variable extinction with $A_V$ up to 8 mag for the Calzetti et al. (2000) attenuation law. We also explored other extinction laws.

The code, initially designed to fit rest-frame UV to near-IR (stellar) emission, can also easily be used to include the thermal mid-IR and beyond. To cover the entire spectral range from the visible to the millimeter domain, and to compare our sources
with SEDs of very different galaxy types, we compiled a great variety of spectral templates from the GRASIL models of Silva et al. (1998), the library of Chary & Elbaz (2001), Rieke et al. (2009), the starburst, Seyfert, and AGN templates of Polletta et al. (2007), the ULIRG templates of Vega et al. (2008), the sub-mm galaxy templates of Michałowski et al. (2010a), and the model fit to SMMJ2135-0102 (Swinbank et al. 2010; Ivison et al. 2010). Extinction can also be added to these spectral templates; SED fits with and without additional extinction will be discussed below.

We carried out both fits of the entire SED (optical, near-IR, and IR) and fits up to 8 μm only (for the Bruzual & Charlot templates). For each template set the free parameters are redshift and (additional) A_V. Physical parameters such as the infrared luminosity, L_{IR}, defined as the luminosity emitted in the range 2–1000 μm; the IR star-formation rate, SFR; and the stellar mass, are subsequently derived from the best-fit templates. In contrast to the IR fits discussed in Sect. 3 we have no handle on the dust temperature, because this is not a parameter describing the SEDs used here. We also checked that the two independent fitting methods used here and in Sect. 3 give consistent results.

### 4.2. Photometric redshifts

As discussed in depth by Laporte et al. (2011), the best-fit photometric redshifts of our sources derived from the optical-to-near-IR photometry (up to 8 μm) and using standard spectral templates is consistently found at z > 7 with a lower probability at low z. This result remains unchanged with the exploration of a wider range of extinction, different attenuation/extinction laws, and template sets used here compared to Laporte et al. (2011) On the other hand, analysis of the IR SED and other arguments clearly favor low redshifts (z ~ 1.5–2.5), as discussed above. Below we therefore limit ourselves to z < 4 and attempt to refine the photometric redshift of the two sources.

### 4.3. Results for z1

Overall the global SED fits for this source are fairly satisfactory, as shown in Fig. 2, albeit with significant discrepancies in the optical domain (cf. below). Several templates (from Polletta et al. 2007; Vega et al. 2008; Michałowski et al. 2010a) also reproduce the 24 μm flux, and the observed 100 and 870 μm fluxes are within 2–3σ of the model. Interestingly, the best-fits for both, the Polletta and Michalowski libraries are found with templates for active galaxies.

The best-fit redshift found with these templates is between z ~ 2.24 and 2.57. The resulting IR luminosity is L_{IR} ~ (2.6–3.2) × 10^{12} L_☉, the corresponding SFR ~ 450–550 M_☉ yr^{-1} using the standard Kennicutt (1998) calibration. Fits to the IR part with the SMMJ2135-0102 template yield z_{phot} ~ 2.0. A somewhat lower redshift of z ~ 1.7 is found with the Bruzual & Charlot (2003) templates using the SED up to the IRAC bands. The estimated extinction is A_V ~ 2.6, the stellar mass M_⋆ ~ 6 × 10^{10} M_☉ for the same Salpeter IMF as adopted by Kennicutt (1998). However, these values should be taken with caution because the fits are poor. For comparison, one obtains M_⋆ ~ 3 × 10^{10} M_☉ from the absolute H-band magnitude (M_B ~ -23.0) using the mass-to-light ratio adopted by Wardlow et al. (2011) for SMGs.

At a more detailed level (see right panel), all spectral templates have some difficulty to reproduce the steep, observed SED between the visible (I, z bands) and the near-IR (Y and J here), and they predict a flux excess in the optical domain. Below we will show that this also holds when variable extinction is added to the empirical templates. The same is also true for all other templates we examined, including the theoretical galaxy

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**Fig. 2.** Fits to the observed SED of source z1 (photometry shown by blue symbols, including 3σ upper limits) using different spectral templates: Chary & Elbaz (2001) (black = best-fit template, and magenta = template with maximum IR luminosity), Polletta et al. (2007) (red = global best-fit template, and green = best-fit to visible-near-IR SED excluding the thermal IR), and Michałowski et al. (2010a) (blue). The best-fit SED with the templates of Vega et al. (2008), very similar to the one using Polletta’s templates, is not shown here for simplicity. Left: global visible to sub-mm SED. Right: zoom on visible to near-IR part of the SED including also the best-fit SED at high-redshift (z = 7.5) from Laporte et al. (2011, dashed line), which is most likely excluded because of our Herschel detections.
templates of Bruzual & Charlot (2003). This sharp drop is of course the reason why this source was selected as an optical dropout (Y-drop).

4.4 Results for Y5

For this source the global fits are less good than for z1. This is because Y5 shows a higher flux ratio between the thermal-IR and the near-IR than z1, whose SED already required templates with extreme IR/near-IR fluxes. For example, the Arp 220 template from Polletta et al. (2007), shown in red, underpredicts the IR flux by a factor \( \gtrsim 5 \). The only template coming close to the observed IR emission is from the SMG library of Michałowski et al. (2010a) (SMMJ221725.97+001238). With a best-fit redshift of \( z_{\text{phot}} \approx 2.15 \) this translates to \( L_{\text{IR}} = 2.2 \times 10^{12} \, L_\odot \), corresponding to \( SFR \approx 380 \, M_\odot \, \text{yr}^{-1} \). Fits to the IR part with the SMMJ2135-0102 template yield \( z_{\text{phot}} \approx 1.8 \). A best-fit redshift of \( z \approx 1.95 \) is found with Bruzual & Charlot (2003) templates using the SED up to the IRAC bands. The estimated extinction is \( A_V \approx 1.6 \), the stellar mass \( M_\star \approx 3 \times 10^{10} \, M_\odot \). However, these values should be taken with caution because the fits in the domain close to the optical are poor. Again, using the absolute H-band magnitude (\( M_H \approx -22.3 \)), one obtains \( M_\star \approx 2 \times 10^{10} \, M_\odot \) with the assumptions already mentioned above.

Similarly to z1, the visible-near-IR drop of the SED (see right panel) is poorly fitted by the spectral templates, predicting that the source should be detectable in the visible (I, \( \gamma \), Y bands in particular), in contrast to our observations. The template that fits this part of the spectrum best is an S0 template from Polletta et al. (2007), shown in green. However, this template underpredicts the IR emission by several orders of magnitudes.

4.5 Possible explanations for the strong SED break/very red spectrum

As already seen, the common, observed spectral templates fail to reproduce the steep, observed SED between the visible (I, \( \gamma \), Y bands and the near-IR (Y and J here), and they predict a flux excess in the optical domain. What causes the sharp observed decrease of the flux between the near-IR and the optical for these sources? The main difficulty arises because the largest spectral break known in galaxy spectra is the Lyman break, whereas the typically observed Balmer (or “4000 Å”) break is smaller than \( z \approx 2.5 \) for the Chary & Elbaz templates, Charlot (2003) templates with the SMC law. Although it is the steepest SED between the optical and the H-band, it falls short in flux in the J-band. For completeness we also examined templates from the synthesis models of Maraston et al. (2006). As expected, these models do not yield significantly different fits in the blue part of the rest-frame optical spectrum.

As already mentioned by Laporte et al. (2011), we also attempted to fit the SEDs with our models including nebular lines (see Schaerer & de Barros 2009). Indeed, in this case the best-fit is found at \( z \approx 1.5 \) such that the [OIII] \( \lambda 4959, 5007 \) lines, and H\beta boost somewhat the J-band flux, and H\alpha the H-band to a lesser extent, contributing thus to the flux decrement between J and Y. However, this solution also requires a very high attenuation (\( A_V \approx 4.0 \) for the SMC law) to reproduce the steeply rising SED toward longer wavelengths\(^4\). Although to the best of our knowledge objects with these red SEDs and strong emission lines are not known, this extreme explanation should be easy to test with spectroscopic observations.

Finally, could composite populations not taken into account by our models help to explain the observed SED? Certainly the theoretical SED models may suffer from this simplification. However, we do not see how this could help to resolve the problem with the large observed spectral break, because a superposition of individual simple stellar populations (not capable of reproducing this observation) can only average out spectral features. We conclude that we have no convincing explanation for the observed sharp drop of the SED of our two sources.

5. Discussion

5.1 Comparison with other objects in the literature

How do our sources compare with other known galaxies and what is their nature?

By design our sources are near-IR-selected, optical dropout sources, i.e., sources with a very red color between the J- and z-band and/or between Y and J. Our sources can therefore be compared to those selected by Capak et al. (2011) from the COSMOS survey. From their Fig. 13 we note that with \( (I - J) > 5.4 \) and \( (z - J) > 3.9 \) and \( 3.1 \) for z1 and Y5, both sources show extreme (very red) optical-to-near-IR colors, when compared to other low-redshift galaxies with red \( (J - z) \) colors. z1 and Y5 are also similar to the z-dropout galaxy HUDF-J2 identified by Mobasher et al. (2005) as a \( z \approx 6.5 \) post-starburst galaxy candidate, but later shown to be most likely at \( z \approx 1.8-2.5 \) interloper by Schaerer et al. (2007); Dunlop et al. (2007); Chary et al. (2007). Although similar in several respects, HUDF-J2 shows a more monotonously rising SED between \( z \) and \( H \) than our objects exhibiting a “sharper” break. The colors and fluxes of our sources are also very similar to those of the lensed optical dropout galaxy #2 found behind the cluster Abell 1835, identified with the \( z = 2.93 \) sub-mm galaxy SMMJ14009+0252 (see Schaerer et al. 2007; Weiß et al. 2009, and references therein).

Our sources are obviously also characterized by a high IR-to-optical flux ratio, a criterion, which has been used by various authors. For example, Rodighiero et al. (2007) have studied IRAC 3.6 \( \mu m \) -selected sources undetected in deep optical HST images. The \( (K-3.6) \) and \( (z-3.6) \) colors of z1 and Y5 are comparable to their sources; the main difference seems to be in \( (H - K) \), where our sources are bluer than those of Rodighiero et al. (2007). z1 and Y5 appear to be related to the very dusty \( z \approx 2-3 \) sources from this study. Fiore et al. (2008) and Dey et al. (2008) have examined 24 \( \mu m \) selected sources with very red colors between 24 \( \mu m \) and the \( R \) band \( (S(24)/S(R) \approx 1000) \). They concluded that the bulk of these sources are very luminous strongly dust-obscured galaxies (referred to as DOGs) at \( z \approx 2 \), powered by AGN and/or by starbursts. Using the I-band

\(^4\) Our model assumes identical attenuation for the continuum and nebular lines, as in Schaerer & de Barros (2009).
as a proxy for R, we obtain a flux ratio \( S(24)/S(R) \gtrsim 20000 \) for z1, an extremely high flux ratio compared to the other samples. From the SED of Y5 (cf. Fig. 3), we also expect this galaxy to show a high 24 μm to optical ratio. According to the source density from Dey et al. (2008), we would have expected \( \sim 4 \) strongly dust-obscured galaxies down to \( \sim 0.3 \) mJy at 24 μm in our \( 45 \) arcmin\(^2 \) field. Because the depth of our MIPS observations is similar, this value is comparable to our source density, although our selection is different. Pope et al. (2008) have also compared DOGs and sub-mm galaxies (SMGs) in the same \( R - K \) colors, showing that \( \sim 30\% \) of SMGs satisfy the DOG criteria, the remainder showing less extreme (i.e. bluer) colors. This confirms that the SEDs of our sources are comparable to a subset of SMGs with the most extreme optical to IR/sub-mm colors, as already seen above (Sect. 4). Our optical data are not deep enough to ascertain whether z1 and Y5 fulfill the usual criteria for extremely red objects (EROs), \( (R - K) \gtrsim 5.6 \) in Vega magnitudes, at least as estimated from \( (I - K) \).

Among known sub-mm galaxies, one source, GN10 or GOODS 850-5, stands out as having particularly extreme IR/sub-mm to visible/near-IR properties, similar to our two galaxies. Indeed, this source is undetected down to \( \sim 0.01 \) μJy (1σ) in the visible, shows fluxes of \( \sim 1-5 \) μJy in the IRAC bands (3.6–8 μm), and peaks at \( \sim 10-20 \) mJy around 1 mm (Wang et al. 2004, 2009; Daddi et al. 2009), quite comparable to z1 and Y5. However, GN10 remains undetected even at JHK (Wang et al. 2009), which can be explained by its higher redshift (\( z \approx 4 \)), recently confirmed from CO spectroscopy (Daddi et al. 2009). The observed spectral break of GN10 found between 3.6 and 2.2 μm and other considerations (Wang et al. 2009; Daddi et al. 2009) suggests a very high attenuation of \( A_V > 4.5-5 \) for this source, or at least for the star-forming part of it, if hosting multiple components. If we assume constant star-formation as Daddi et al. (2009) for their SED modeling, we would infer \( A_V \approx 3 \) (7.8) mag for z1 (Y5). Comparing the infrared-derived SFR with the upper limits in the rest-frame UV domain, we can also estimate the attenuation of our sources. Adopting the f-band flux as a constraint for the UV flux at \( \sim 2300-2500 \) Å and using the Kennicutt (1998) calibration, we obtain \( A_V \approx 4 \) mag for both sources.

Sub-mm galaxies are also known to exhibit very strong attenuation. For example, sources with Balmer decrement measurements indicate \( A_V \sim 1-3 \) (Swinbank et al. 2004; Takata et al. 2006), and from SED fits (Swinbank et al. 2004) estimate \( A_V = 3.0 \pm 1.0 \) for their sample. Wardlow et al. (2011) find \( A_V = 2.6 \pm 0.2 \) from the median SED of sub-mm galaxies, but more extreme attenuations are found within the sub-mm
5.2. The other high-z candidates of the survey

Based on our FIR detections we have identified two potential interlopers among the ten high-z candidates discovered by Laporte et al. (2011). Most of the other candidates are in crowded regions where some sources emit in the FIR and are blended with each other, making any FIR measurement impossible. Two other candidates only seem to be clean from any contamination in the MIPS, PACS and SPIRE maps, namely Y3 and Y4. They remain undetected in all bands. However, the FIR upper limits obtained do not allow us to distinguish between low and high-redshift. On the other hand we can rule out that these sources are as extreme as \(z = 7\) and \(z = 5\) in their IR/sub-mm to near-IR flux ratio, because they should otherwise clearly be detected in our Herschel images.

6. Conclusions

Analyzing the FIR SED of two high-redshift dropout candidates we find that both galaxies are likely at \(z \sim 2\) rather than \(z > 7\). From the FIR point of view alone, both galaxies could be similar to ULIRGs or SMGs, which are common at \(z \sim 2\). At \(z > 7\) the SEDs would imply extreme dust temperatures and luminosities. Fitting the global SEDs considering all available data from the near-IR and in the IR, than the typical sub-mm galaxies (cf. Wardlow et al. 2011). Finally, our sources stand out by their large spectral break, which — to the best of our knowledge — is unusual among intermediate-redshift sources.