From star-forming spirals to passive spheroids: integral field spectroscopy of E+A galaxies

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ABSTRACT
We present three-dimensional spectroscopy of 11 E+A galaxies at \( z = 0.06–0.12 \). These galaxies were selected for their strong H\( \delta \) absorption but weak (or non-existent) [O II] \( \lambda 3727 \) and H\( \alpha \) emission. This selection suggests that a recent burst of star formation was triggered but subsequently abruptly ended. We probe the spatial and spectral properties of both the young (\( \lesssim 1 \) Gyr) and old (\( \gtrsim \) few Gyr) stellar populations. Using the H\( \delta \) equivalent widths we estimate that the burst masses must have been at least 10 per cent by mass (\( M_{\text{burst}} \gtrsim 10^{10} M_\odot \)), which is also consistent with the star formation history inferred from the broad-band spectral energy distributions. On average the A stars cover \( \sim 33 \) per cent of the galaxy image, extending over 2–15 kpc\(^2\), indicating that the characteristic E+A signature is a property of the galaxy as a whole and not due to a heterogeneous mixture of populations. In approximately half of the sample, we find that the A stars, nebular emission and continuum emission are not co-located, suggesting that the newest stars are forming in a different place than those that formed \( \lesssim 1 \) Gyr ago, and that recent star formation has occurred in regions distinct from the oldest stellar populations. At least 10 of the galaxies (91 per cent) have dynamics that class them as ‘fast rotators’ with magnitudes, \( v/\sigma, \lambda_B \) and bulge-to-total (B/T) ratio comparable to local, representative ellipticals and S0s. We also find a correlation between the spatial extent of the A stars and the dynamical state of the galaxy such that the fastest rotators tend to have the most compact A star populations, providing new constraints on models that aim to explain the transformation of later type galaxies into early types. Finally, we show that there are no obvious differences between the line extents and kinematics of E+A galaxies detected in the radio (active galactic nucleus, AGN) compared to non-radio sources, suggesting that AGN feedback does not play a dramatic role in defining their properties, and/or that its effects are short.

Key words: galaxies: evolution – galaxies: kinematics and dynamics – galaxies: starburst – galaxies: stellar content.

1 INTRODUCTION
Galaxies with strong Balmer absorption lines in their spectra, but weak nebular emission (such as [O ii] \( \lambda 3727 \) Å), represent a short-lived but potentially important phase in galaxy evolution (e.g. Tran et al. 2003, 2004). These ‘E+A’ galaxies have strong absorption lines (such as H\( \delta \)), representing a stellar population dominated by A stars, which are either absent or overwhelmed by the much brighter OB stars in most galaxies. These signatures suggest that the star formation within the galaxy abruptly ended \( \lesssim 1 \) Gyr ago, possibly following a starburst phase. It is likely that a variety of physical mechanisms lead to such a stellar population (e.g. Dressler & Gunn 1982; Couch & Sharples 1987; Zabludoff et al. 1996; Balogh et al. 1999; Poggianti et al. 1999), but most invoke a major transformation from one galaxy type to another, possibly representing an evolutionary link between gas-rich, star-forming galaxies and quiescent spheroids. Although such galaxies are very rare in the local Universe, their short lifetime means they could potentially represent an important phase in the evolution of most normal galaxies (e.g. Zabludoff et al. 1996; Yan et al. 2009).

In order to trace the route by which star-forming galaxies evolve into quiescent systems, there are several issues which must be addressed. First, from where in the galaxy does the unusual spectrum
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1. Introduction

2. Observations and Data Reduction

3. Results

4. Discussion

5. Conclusion

Table 1. Target information. Coordinates, redshifts and exposure times of the galaxies in our sample. The two E+A galaxies labelled "++" (J0906+5221 and J1242+0237) are both in cluster environments. Note that J2013+0116 was selected to have weak emission lines, so is an e(α), rather than E+A, galaxy.
to ensure reliable dynamical measurements could be made. Thus, it has the strongest \([\text{O} \text{II}]\) emission in our sample \((W_{\text{\textsc{[O}ii]}} = 5.2 \pm 0.3 \text{ \AA})\), and is more properly classified as an e(a) galaxy (Poggianti et al. 1999). This is therefore a fairly heterogeneous sample, by design.

In Fig. 1 we show where these galaxies lie in the luminosity–linewidth plane of the Faber–Jackson relation (Faber & Jackson 1976). We measure the velocity dispersions from the SDSS spectra, using the cross-correlation technique described in Section 2.2. We use the linewidth of the Mg\,\text{i} triplet as a proxy for the stellar kinematics of an old stellar population. The E+A galaxies in our sample span a range of velocity dispersions from \(\sigma \sim 80\) to \(180 \text{ km s}^{-1}\), and nearly two magnitudes in the \(K\) band. For comparison we also plot the early-type galaxies selected from SDSS (Bernardi et al. 2005), cross-matched to Two Micron All Sky Survey (2MASS). To ensure a fair comparison with the E+A galaxies, we limit the redshift range of the comparison sample to \(z < 0.1\), and apply no \(k\)-corrections to either sample. As Fig. 1 shows, the E+A galaxies in our sample are systematically brighter in the \(K\) band for a fixed velocity dispersion \((\Delta M_K = 0.6 \pm 0.2)\). A comparable offset in \(R\) band \((\Delta m_R \sim 0.6 \text{ mag})\) was also noted by Norton et al. (2001) who used long-slit spectroscopy of 20 E+A galaxies from the sample described by Zabludoff et al. (1996).

The \((u-g)\) versus \((r-i)\) colours of the galaxies in our sample are shown in Fig. 2. We compare these to the galaxies from SDSS Data Release 5 (DR5) with redshift range \(0.05 < z < 0.1\) and magnitude range \(16 < r < 17\). We also compute three evolutionary \(r\)-tracks: (i) an exponentially declining star formation rate (SFR), with \(\tau = 4\) Gyr, and \(r_v = 1\) mag extinction; (ii) a ‘truncated’ model which shows the colour evolution over a 2-Gyr period following truncation of star formation in the model; (iii) a model in which a 10 per cent (by mass) instantaneous burst is superposed upon the old stellar population (filled circle); the evolution from blue to red is followed for 2 Gyr following the burst. This latter model (which includes \(r_v = 1\) mag extinction) provides a reasonable match to the colours of our E+A galaxies, as shown in Balogh et al. (2005).

Finally, we note that of 10 galaxies in our sample that are covered as part of the VLA Faint Images of the Radio Sky at Twenty-centimetre (FIRST) survey, two are detected at \(>5\sigma\) (J0948+0230 and J1013+0116) and two more are detected at \(3\sigma–5\sigma\) (J0835+4239 and J1642+4153). These detections have 1.4-GHz fluxes in the range \(0.5–2.2\) mJy, corresponding to luminosities of \(L_{1.4} = 8–15 \times 10^{23} \text{ W Hz}^{-1}\). We return to this in Section 4.

2.2 GMOS spectroscopic imaging

Spectro-imaging observations of 10 new E+A galaxies in our sample were taken with the GMOS-North and GMOS-South IFU between 2005 and 2007 (Allington-Smith et al. 2002). Seven targets were observed in ‘stare’ mode using one-slit mode which results in a field of view of \(5 \times 7\) arcsec\(^2\), while four targets were observed with nod-and-shuffle in two-slit mode (resulting in a field of \(5 \times 5\) arcsec\(^2\)). For observations taken with nod-and-shuffle we chopped away from the target by 30 arcsec every 30 s (see Swinbank et al. 2005, for a detailed discussion of the observing procedure). The exposure times for each target were typically 11–18 ks, with longer exposures taken on fainter targets (Table 1). All observations were carried out in dark time, typically with 0.6-arcsec seeing in \(V\) band.
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3 ANALYSIS

3.1 Stellar masses and morphologies

Before discussing the spatially resolved properties, we first make use of the existing imaging to estimate the stellar masses and morphologies of the galaxies in our sample. To estimate the stellar masses, we follow McGee et al. (2011) and use spectral energy distribution modelling employing the Sloan $ugriz$- and $K$-band photometry (Balogh et al. 2005). Briefly, we generate galaxy templates using Bruzual & Charlot (2003) models with a Salpeter initial mass function (IMF; Salpeter 1955) assuming a lower and upper IMF mass cut-off of 0.1 and 100 $M_\odot$, respectively. We follow the galaxy parameter ranges used by Salim et al. (2007). The model spectra include a range of galaxy age (0.1–13 Gyr), metallicity (0.005–2.5 $Z_\odot$), star formation history and dust obscuration ($A_v = 0$–6 mag). In particular, the exponentially declining SFRs with superimposed bursts are randomly chosen with a uniform distribution, and we allow bursts that last some time randomly distributed in duration between 30 and 300 Myr. The strength of the bursts are also randomly chosen so that during the lifetime of the burst they produce between 0.03 and 4 times the stellar mass the galaxy had at the onset of the burst. From these model star formation histories, we generate magnitudes by convolving the resulting spectra with the $ugriz$ and $K$-band filter response curves. To fit to the observations, we search the entire parameter space and minimize the $\chi^2$, but generate a probability distribution function for each parameter for each galaxy (see McGee et al. 2011, for a detailed discussion). The resulting stellar masses (and their $\sigma$ errors) are given in Table 2. The median stellar mass of the galaxies in our sample is $8 \pm 2 \times 10^{10} M_\odot$.

Since the $K$-band luminosity is a good tracer of the stellar mass (and is less sensitive to the recent star formation history than the optical bands), we can also perform a simple check of the stellar mass using a canonical mass-to-light ratio. Balogh et al. (2005) (see also Fig. 2) show that the $(r - g)$ and $(r - i)$ colours of $\text{E}\!+\!\text{A}$ galaxies are consistent with a truncated star formation model in which a burst of 5–15 per cent is superimposed on to an old population, and use this to derive a canonical $M/L_K = 0.8 \pm 0.1$. Applying this to our $K$-band magnitudes, we derive $M_\star = 9 \pm 2 \times 10^{10} M_\odot$, which is consistent with the estimates from the more sophisticated modelling above. This stellar mass is also comparable to the average stellar mass of the parent sample from Balogh et al. (2005): $M_\star = 8 \pm 2 \times 10^{10} M_\odot$.

The SDSS imaging is also useful for measuring Sérsic indices, asymmetries and radial colour gradients within the galaxies which, along with the distribution of $H\delta$ equivalent widths, may provide a diagnostic of the physical mechanism(s) of $\text{E}\!+\!\text{A}$ formation. For example, positive colour gradients (i.e. bluer in centre) can arise when young stars are more concentrated than the old stellar population. In this case, the distribution of $H\delta$ equivalent widths are also likely to be compact. In contrast, a negative colour gradient, together with a positive $H\delta$ equivalent width radial gradient, may represent a galaxy in which the molecular clouds are not confined to the nuclear regions.

Table 2. Properties of the $\text{E}\!+\!\text{A}$ galaxies in our sample.

<table>
<thead>
<tr>
<th>Equivalent width</th>
<th>$H\delta$</th>
<th>B/T</th>
<th>Sérsic index</th>
<th>$M_\star$</th>
<th>$[\text{O} \text{n}]$ flux</th>
<th>SFR($[\text{O} \text{n}]$)</th>
<th>$v \sin i/\sigma$</th>
<th>$\text{inc}$</th>
<th>$v_c$</th>
<th>$\lambda_R$</th>
<th>$r(H\delta)$</th>
<th>$A_{H\delta} &gt; 6$</th>
</tr>
</thead>
<tbody>
<tr>
<td>[\text{O} \text{n}]</td>
<td>(Å)</td>
<td>(Å)</td>
<td>(n)</td>
<td>($\times 10^{10}$)</td>
<td>($\times 10^{-16}$ erg s$^{-1}$ cm$^{-2}$)</td>
<td>(M$\odot$ yr$^{-1}$)</td>
<td>(°)</td>
<td>(km s$^{-1}$)</td>
<td>(cm$^{-1}$)</td>
<td>(MPc)</td>
<td>(km s$^{-1}$)</td>
<td></td>
</tr>
<tr>
<td>J0835+4239</td>
<td>2.9 ± 0.2</td>
<td>6.9 ± 0.3</td>
<td>0.48</td>
<td>7 ± 1</td>
<td>8.3$^{+0.8}_{-0.6}$</td>
<td>5.0 ± 0.5</td>
<td>0.14 ± 0.07</td>
<td>0.4 ± 0.1</td>
<td>51 ± 11</td>
<td>55 ± 15</td>
<td>0.11 ± 0.04</td>
<td>2.4 ± 0.2</td>
</tr>
<tr>
<td>J0906+5221*</td>
<td>1.5 ± 0.3</td>
<td>6.3 ± 0.4</td>
<td>0.69</td>
<td>4 ± 1</td>
<td>2.0$^{+0.5}_{-0.3}$</td>
<td>2.3 ± 1.0</td>
<td>0.08 ± 0.04</td>
<td>1.0 ± 0.1</td>
<td>26 ± 5</td>
<td>75 ± 12</td>
<td>0.27 ± 0.04</td>
<td>2.0 ± 0.2</td>
</tr>
<tr>
<td>J0938+0001</td>
<td>3.9 ± 0.4</td>
<td>5.0 ± 0.3</td>
<td>0.47</td>
<td>3.5 ± 1.0</td>
<td>23.4 ± 6.0</td>
<td>5.1 ± 0.5</td>
<td>0.15 ± 0.04</td>
<td>5.5 ± 0.1</td>
<td>1.5 ± 0.6</td>
<td>3.5 ± 1.1</td>
<td>0.3 ± 0.01</td>
<td>2.1 ± 0.01</td>
</tr>
<tr>
<td>J0948+0230</td>
<td>2.4 ± 0.5</td>
<td>6.4 ± 0.2</td>
<td>0.29</td>
<td>2.2 ± 0.3</td>
<td>6.6$^{+0.4}_{-0.3}$</td>
<td>3.8 ± 0.9</td>
<td>0.05 ± 0.03</td>
<td>1.8 ± 0.1</td>
<td>54 ± 7</td>
<td>200 ± 10</td>
<td>0.53 ± 0.03</td>
<td>1.2 ± 0.04</td>
</tr>
<tr>
<td>J1013+0116</td>
<td>5.2 ± 0.3</td>
<td>6.8 ± 0.5</td>
<td>0.46</td>
<td>&gt;7</td>
<td>15.0$^{+0.2}_{-0.1}$</td>
<td>12.1 ± 0.5</td>
<td>0.48 ± 0.12</td>
<td>0.2 ± 0.1</td>
<td>42 ± 10</td>
<td>152 ± 14</td>
<td>0.27 ± 0.04</td>
<td>2.8 ± 0.3</td>
</tr>
<tr>
<td>J1242+0237*</td>
<td>2.0 ± 0.4</td>
<td>6.0 ± 0.2</td>
<td>0.40</td>
<td>3.8 ± 0.5</td>
<td>1.9$^{+0.5}_{-0.2}$</td>
<td>2.7 ± 0.6</td>
<td>0.06 ± 0.03</td>
<td>0.5 ± 0.1</td>
<td>21 ± 5</td>
<td>92 ± 6</td>
<td>0.20 ± 0.03</td>
<td>1.0 ± 0.2</td>
</tr>
<tr>
<td>J1348+0204</td>
<td>2.1 ± 0.5</td>
<td>6.0 ± 0.3</td>
<td>0.67</td>
<td>6.2 ± 0.4</td>
<td>6.9$^{+0.4}_{-0.3}$</td>
<td>7.1 ± 0.8</td>
<td>0.11 ± 0.03</td>
<td>0.4 ± 0.1</td>
<td>49 ± 7</td>
<td>63 ± 8</td>
<td>0.10 ± 0.04</td>
<td>1.6 ± 0.2</td>
</tr>
<tr>
<td>J1627+4800</td>
<td>0.8 ± 0.4</td>
<td>5.2 ± 0.5</td>
<td>0.73</td>
<td>&gt;7</td>
<td>5.7$^{+0.3}_{-0.2}$</td>
<td>&lt;2.1</td>
<td>&lt;0.13</td>
<td>2.2 ± 0.1</td>
<td>37 ± 5</td>
<td>229 ± 12</td>
<td>0.52 ± 0.05</td>
<td>2.1 ± 0.5</td>
</tr>
<tr>
<td>J1642+4153</td>
<td>1.2 ± 0.5</td>
<td>6.5 ± 0.5</td>
<td>0.64</td>
<td>3.5 ± 0.6</td>
<td>6.0$^{+0.1}_{-0.1}$</td>
<td>4.5 ± 1.0</td>
<td>0.08 ± 0.03</td>
<td>1.2 ± 0.1</td>
<td>51 ± 5</td>
<td>201 ± 10</td>
<td>0.41 ± 0.04</td>
<td>1.5 ± 0.3</td>
</tr>
<tr>
<td>J1715+5822</td>
<td>2.2 ± 0.5</td>
<td>6.4 ± 0.8</td>
<td>0.48</td>
<td>4.5 ± 1.0</td>
<td>7.9$^{+0.4}_{-0.3}$</td>
<td>2.8 ± 0.5</td>
<td>0.17 ± 0.05</td>
<td>0.4 ± 0.1</td>
<td>49 ± 5</td>
<td>90 ± 5</td>
<td>0.46 ± 0.03</td>
<td>2.7 ± 0.3</td>
</tr>
<tr>
<td>J2307+1525</td>
<td>1.7 ± 0.3</td>
<td>6.3 ± 0.4</td>
<td>0.45</td>
<td>&gt;7</td>
<td>4.5$^{+0.3}_{-0.2}$</td>
<td>3.0 ± 0.5</td>
<td>0.05 ± 0.02</td>
<td>0.7 ± 0.1</td>
<td>–</td>
<td>–</td>
<td>0.22 ± 0.03</td>
<td>1.4 ± 0.2</td>
</tr>
</tbody>
</table>

Notes. The two $\text{E}\!+\!\text{A}$ galaxies labelled * (J0906+5221 and J1242+0237) are both in cluster environments. $n$ denotes the Sérsic index measured from the SDSS $gri$ photometry. The estimate of the SFR is calculated using the $[\text{O} \text{n}].3727$ emission line flux and assuming the calibration from Kennicutt (1998).
From our sample of 11 galaxies, three show positive colour gradients \[ \Delta (g - i)/\Delta \log r > 0.0 \], and the rest are either flat or mildly negative (see also Fig. 5 and Yang et al. 2008). The galaxies with bluer nuclei (positive colour gradients) may evolve into the negative gradients typical in E/S0s if the central parts of these galaxies are metal-enhanced (Yang et al. 2008). Finally, we note that the median asymmetry and Sérsic index of the sample is 0.03 ± 0.01 and 4.9 ± 1.3 respectively, characteristic of bulge-dominated systems.

3.2 Spatial distributions

To investigate the spatial distribution of the gas and stars in the galaxies, first we extract narrow-band slices from the data cube around the emission and absorption lines of interest. For a given emission or absorption line, we fit and subtract the continuum using a region ± 100 Å from the centre of the emission or absorption feature using a 3σ clip in the fit to be sure that neighbouring emission and absorption lines are omitted. We also extract continuum images from the data cube. For each image, we median filter each spectral pixel in the data cube between 4350 and 4750 Å (rest frame) as a proxy for the ‘old’ stellar population (above the 4000-Å break) and between 3650 and 3850 Å (rest frame) as a proxy for the ‘young’ stellar population and show these in the second row of Fig. 3. We caution that the continuum emission from above and below the 4000-Å break has contributions from both ‘old’ and ‘young’ stars. For example, for a 10 per cent by mass starburst which has been truncated at 1 Gyr, for solar metallicity the contributions to the continuum emission from the burst to the total 3650–3850 and 4350–4750 Å continuum emission are ~60 and 40 per cent respectively, but can vary from 40 to 70 per cent in each band depending on starburst age and metallicity (Maraston 1998).

A better test of how the A stars are related to the underlying stellar emission (as well as the current star formation activity) is to turn to the Hδ absorption line. We estimate the equivalent width of Hδ at each spatial pixel following Balogh et al. (1999) (see also Goto et al. 2003). We first average each spatial pixel over a 0.6 × 0.6 arcsec^2 region to approximately match the seeing, and estimate the continuum flux using 3σ clipped linear interpolation between two wavelength windows placed at either side of the Hδ line (4030–4082 and 4122–4170 Å; Goto et al. 2003). This 3 × 3 averaging produces an effective seeing disc of 0.85 arcsec, and this is taken into account in all of the following calculations. The rest-frame equivalent width of the Hδ absorption line is then calculated by summing the ratio of the flux in each pixel of the spectrum, over the estimated continuum flux in that pixel based on our linear interpolation; its distribution over the area of each galaxy is shown in Fig. 4. In this figure we also indicate the centroids of the nebular emission and continuum from above the 4000-Å break, with a 1σ error ellipse. To calculate these error ellipses we compute the number of photons per pixel with their associated \( \sqrt{n} \) uncertainties, and recompute centroid with 10^5 Monte Carlo realizations. It is interesting to note that the regions of strongest Hδ absorption do not always coincide.

Figure 3. Results from the IFU observations of the E+A galaxies in our sample. From top to bottom: (1) true colour BIK-band images of each of the galaxies constructed from Gemini/GMOS imaging (blue), SDSS r-band (green) and UKIRT/UFTI K-band imaging (red). In each panel we overlay the GMOS IFU field of view. The arrows denote the directions of north and east. (2) Continuum images from the continuum above the 4000-Å break, derived by collapsing the data cube between 4350 and 4750 Å in the rest frame, with dark grey indicating high intensity. The contours trace the continuum from the rest-frame 3650–3850 Å continuum emission (which should have a stronger contribution from young stars than the continuum emission above the 4000-Å break). The solid bar in each panel denotes the seeing disc for the observations. (3) Gas phase emission line intensity (derived from either [O II] λ3727 or [O III] λ5007, as indicated). The contours start at 3σ and are incremented by 1σ. (4) Two-dimensional stellar velocity fields of the galaxies, measured from the Hδ line. (5) Two-dimensional stellar velocity dispersion maps. The dotted line marks the major kinematic axis in the velocity fields, where it can be clearly identified.
Figure 4. The distribution of rest-frame H$\delta$ equivalent widths across each of the galaxies in our sample. The contours are separated by $\Delta W_o (H\delta) = 1\, \AA$. The ‘X’ denotes the centre of the galaxy as defined from the centre of the continuum above the 4000-Å break, while the ellipses denote the centre and 1σ uncertainty in the nebular emission centroid (red) and continuum emission centroid (black). This figure highlights that six E+A galaxies in our sample have H$\delta$ absorption, nebular emission and continuum emission which are spatially co-located on ∼kpc scales (J0948, J1642, J1627, J1715, J2307), whilst three systems have significant offsets between nebular emission and H$\delta$ (J0906, J1242, J1348). Two galaxies have nebular emission and H$\delta$ in agreement, but with very extended H$\delta$ and nebular emission which is offset from the strongest continuum (J0835, J1013), suggesting that the A and OB stars formed in a different location than the older stars. At least in some cases, the new OB stars are either displaced from the older A stars. This reconciles the fact that both emission and strong absorption can be seen in the same, integrated galaxy spectrum – the two populations do not generally arise from exactly the same place.
for the galaxies where nebular emission and H$\delta$ emission are separated by the E and A stars. On average, the centroid of the A stars and the continuum A-star/gas/continuum emission, whilst half are clearly more complex. Thus, approximately half the sample are (relatively) simple in their dynamics as measured from $v\sin(i)/\sigma$ for the E+A galaxies which are pressure supported; the average $v\sin(i)/\sigma$ for galaxies with $A$(H$\delta$)$>6$ Å is 0.5 ± 0.2, compared with $v\sin(i)/\sigma = 1.1 ± 0.5$ for the others. The spatial extent of the A stars is sensitive to detecting the continuum, we normalize the area, $A$(H$\delta$), by the spatial extent of the galaxy, as measured from the continuum above the 4000-Å break. We note that the results are insensitive to whether we use the SDSS $r$-band image to measure the half-light radius or $R_e$. This plot shows that the galaxies with the most compact H$\delta$ distribution tend to be those which have dynamics consistent with rotation, while the galaxies with A stars which are most widespread are those with dispersion-dominated dynamics. Bottom-right panel: the spatial extent of the A-star population normalized to the continuum as a function of $\lambda_{\kappa}$ (Emsellem et al. 2007). High and low $\lambda_{\kappa}$ correspond to ‘fast’ and ‘slow’ rotators, respectively. This plot shows that the fastest rotators tend to have the most compact A-star spatial distribution. We note that although our sample is small, the galaxies with weak AGN do not stand out in terms of their A-star extent or dynamics from the rest of the sample in any of the panels.

with regions with the strongest emission lines. Indeed, Fig. 4 shows that five objects have H$\delta$, nebular emission and continuum emission which are spatially co-located on ∼kpc scales (J0948, J1642, J1627, J1715, J2307), whilst three systems have significant offsets between nebular emission and H$\delta$ (J0906, J1242, J1348). Two galaxies have nebular emission and H$\delta$ in agreement, but with very extended H$\delta$ which is offset from the strongest continuum (J0835, J1013). Thus, approximately half the sample are (relatively) simple in their A-star/gas/continuum emission, whilst half are clearly more complex. On average, the centroid of the A stars and the continuum emission are separated by $\Delta r_{\text{H}\delta-\text{gas}} = 0.8 ± 0.4$ kpc, whilst the average offset between the H$\delta$ and nebular emission is $\Delta r_{\text{H}\delta-\text{gas}} = 1.0 ± 0.2$ kpc, but can be as large as 1.7 kpc. This suggests that at least in some cases, the new OB stars are either displaced from the older A stars, or that the extinction is not uniform. In either case, this helps to reconcile the fact that both emission and strong absorption can be seen in the same, integrated galaxy spectrum. The two populations do not generally arise from exactly the same place.

Next, we examine the spatial extent of the A stars through the H$\delta$ absorption. In the following calculations, we assume that the continuum emission is smoothly distributed and subtract the amplitude of the PSF at the redshift of the galaxy in quadrature from the areas covered by the continuum and A stars. As shown by the distributions in Fig. 4, the A stars are typically distributed over a large radius. We quantify this by computing the radius at which the H$\delta$ equivalent width drops to half the peak value (this is extracted along the major kinematic axes within the galaxy which is discussed in Section 3.3). This quantity, $r_{\text{h}}(W, H\delta)$, is compared with the observed $g - i$ colour gradient within the central 2 kpc in the top-left panel of Fig. 5. Positive radial colour gradients

Figure 5. The spatial extent of the A stars within the E+A galaxies in our sample as a function of colour gradient and kinematics. Top-left panel: the extent of the A stars (defined as the radius at which the H$\delta$ equivalent width drops to half the peak value along the major kinematic axis) as a function of the $g - i$ colour gradient within the central 2 kpc of the galaxy. Positive radial colour gradients denote light that is bluer in the centre. The dashed line denotes a colour gradient of zero, and we use this to differentiate galaxies with positive (black filled circles) from negative colour gradients (red filled squares). In all panels, we identify the E+A galaxies which have weak radio emission ($L_{1.4} < 10^{22}$ W Hz$^{-1}$, highlighted with an open circle). Top-right panel: the spatial extent of the A stars as a function of the dynamics as measured from $v\sin(i)/\sigma$ for the E+A galaxies in our sample. We also include the six E+A galaxies from Pracy et al. (2009). To provide a quantitative measure of the spatial extent of the A stars, we define $A$(H$\delta$) as the area over which the H$\delta$ is stronger than 6 Å. The galaxies with higher $v\sin(i)/\sigma$ tentatively show signs of having more compact A-star distributions compared to galaxies which are pressure supported: the average $v\sin(i)/\sigma$ for the galaxies where $A$(H$\delta$)$>6$ Å is 0.5 ± 0.2, compared with $v\sin(i)/\sigma = 1.1 ± 0.5$ for the others. Bottom-left panel: since the measurement of the spatial extent of H$\delta$ is sensitive to detecting the continuum, we normalize the area, $A$(H$\delta$), by the spatial extent of the galaxy, as measured from the continuum above the 4000-Å break. We note that the results are insensitive to whether we use the SDSS $r$-band image to measure the half-light radius or $R_e$. This plot shows that the galaxies with the most compact H$\delta$ distribution tend to be those which have dynamics consistent with rotation, while the galaxies with A stars which are most widespread are those with dispersion-dominated dynamics. Bottom-right panel: the spatial extent of the A-star population normalized to the continuum as a function of $\lambda_{\kappa}$ (Emsellem et al. 2007). High and low $\lambda_{\kappa}$ correspond to ‘fast’ and ‘slow’ rotators, respectively. This plot shows that the fastest rotators tend to have the most compact A-star spatial distribution. We note that although our sample is small, the galaxies with weak AGN do not stand out in terms of their A-star extent or dynamics from the rest of the sample in any of the panels.

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indicate galaxies that are bluer in the centre, and are measured in only three galaxies (∼30 per cent). This is lower than the fraction of E+A galaxies with strong positive colour gradients (i.e. blue cores) found by Yamauchi & Goto (2005) (∼67 per cent), but still higher than the fraction of early-type galaxies with positive colour gradients (Balcells & Peletier 1994). Within our sample, there is no evidence for the correlation between colour gradient and spatial extent of the A stars (as might have been expected from merger or tidal models for E+A formation; Bekki et al. 2005).

Another way to characterize the extent of the Hβ absorption feature is to measure the area over which the equivalent width exceeds some threshold. We choose a threshold of 6 Å and calculate the physical area A(Hβ > 6 Å) for each galaxy. The results are shown in Fig. 5 as a function of various kinematic parameters, discussed below. Note that in all but one galaxy (J0948+0230) the area with \( W_o(\text{H}\beta) > 6 \) Å is significantly larger than the seeing. The median A(\( \text{H}\beta > 6 \) Å) = 4.0 ± 2.1 kpc\(^2\), with a range of A(\( \text{H}\beta > 6 \) Å) = 0.8–15.2 kpc\(^2\). In comparison to the continuum emission from above the 4000 Å break, the A stars cover at least ∼10 per cent of the area, with a median of 33 ± 5 per cent. Such widespread A-star populations have also been found in previous samples: Goto et al. (2008) show that the A stars in two SDSS-selected E+A galaxies are distributed over ∼10 kpc\(^2\), while the seven E+A galaxies from Pracy et al. (2009) have A stars that are typically distributed over 5–15 kpc\(^2\).

In summary, we note two important observations. First, the Hβ is widespread, typically over A(\( \text{H}\beta > 6 \) Å) = 4.0 ± 2.1 kpc\(^2\) and thus the unusual spectrum characteristic of E+A galaxies is a property of the galaxy as a whole. It is not due to a heterogeneous mixture of populations, such as a nuclear starburst with a very high equivalent width (≥10 Å) on top of an otherwise normal galaxy. Secondly, in approximately half of our sample, the A stars, nebular emission and continuum are not co-located, suggesting that the newest stars are often forming in a different place than those that formed ≤1 Gyr ago, and that recent star formation has occurred in regions distinct from the continuum emission (the older stellar populations).

### 3.3 Kinematics

To obtain the stellar velocity structure within the galaxies, we first construct a continuum image of the galaxy from the data cube collapsed between 4350 and 4750 Å, and bin the data to a constant signal-to-noise ratio per bin using the Voronoi binning method (Cappellari & Copin 2003). We demand a signal-to-noise ratio of 30 in continuum, and then cross-correlate each of the individual spatial regions with a library of template spectra (Vazdekis & Ari-moto 1999). The cross-correlation is performed using the penalized pixel fitting method (PPXF; Cappellari & Emsellem 2004) and we demand a signal-to-noise ratio of 6 Å and calculate the one-dimensional rotation curve and velocity dispersion profile along and show these in Fig. 6. From each of the two-dimensional maps, we also extract the distribution of Hβ equivalent widths, both in radial bins and along the kinematic cross-section.

A measure of whether the dynamics are dominated by ordered- or random- motion can be obtained by measuring the ratio of rotational velocity to line-of-sight velocity dispersion, \( v / \sigma \). We derive a median luminosity-weighted \( v / \sigma = 0.8 ± 0.4 \) for the whole sample with a range of 0.2–5.5. However, Fig. 7 shows that there is a weak correlation between absolute magnitude and \( v / \sigma \), such that the lower luminosity systems have an increasing rotational support, although only four of the 11 galaxies have \( v / \sigma > 1 \), uncorrected for inclination. Recent studies of early-type galaxies from Emsellem et al. (2007) identify a comparably weak trend of \( v / \sigma \) with magnitude (although we note that their sample has a slightly lower average \( v / \sigma \) with an average \( v / \sigma = 0.30 ± 0.03 \)). Indeed, the early types in Emsellem et al. (2007) have \( M_v = -21.8 ± 0.3 \) compared to \( M_v = -21.5 ± 0.5 \) for the E+A galaxies, and with stellar masses of \( 1.0 ± 0.3 \times 10^{11} \) M⊙, which is also consistent with our sample. Thus, these early types and ellipticals from Emsellem et al. (2007) also have many of the properties expected for the descendant population of E+A galaxies.

Bender, Saglia & Gerhard (1994) also find a correlation between magnitude and \( v / \sigma \) for luminous elliptical galaxies. Although the Bender et al. (1994) sample is dominated by brighter galaxies than the E+A galaxies studied here (typically ∼1.3 mag brighter), those galaxies with high \( v / \sigma \) (∼1 and so comparable to the E+A galaxies) are increasing dominated by lower mass, lower luminosity Sb–S0s.

A better test of the dynamical state of galaxies with low \( v / \sigma \) can be made by measuring the luminosity-weighted stellar angular momentum per unit mass, \( \lambda_R \), as developed by Emsellem et al. (2007). This combines the velocity amplitude and line-of-sight velocity dispersion: \( \lambda_R = (R/v)|/(R/\sqrt{v^2 + \sigma^2}) \), where R, v and \( \sigma \) are the radius, rotational velocity and velocity dispersion from the kinematic centre. \( \lambda_R \) is designed to quantitatively distinguish galaxies that have similar \( v / \sigma \) values but very different velocity structures by identifying those galaxies with significant angular momentum per unit mass. Emsellem et al. (2007) define ‘fast’ and ‘slow’ rotators as \( \lambda_R > 0.1 \) and \( \lambda_R < 0.1 \), respectively. Classic ‘fast’ rotators are dominated by low-mass spheroids (with a large \( v / \sigma \) whilst ‘slow’ rotators are dominated by ellipticals, albeit with a large range in \( M_R \). We report \( \lambda_R \) for each galaxy in Table 2. Our sample has a median \( \lambda_R = 0.35 ± 0.1 \), and all of the galaxies in our sample have \( \lambda_R > 0.1 \). When including the E+A galaxies from Pracy et al. (2009), 18/19 galaxies with well-resolved dynamical maps have \( \lambda_R > 0.1 \). Taken with the measurements of \( v / \sigma , \lambda_R \) and stellar masses, the E+A galaxies in our sample have properties comparable to those measured in local, representative elliptical galaxies and S0s where similar measurements have been made.

To investigate how the dynamics and A star populations are related, we correlate the kinematics as measured from \( v / \sigma \) and \( \lambda_R \) as
Figure 6. Extracted, one-dimensional equivalent width, velocity and line-of-sight velocity dispersion profiles are shown for each of the E+A galaxies in our sample. The profiles are extracted along the major axis as indicated in Fig. 3. Left: one-dimensional extent of equivalent widths. Middle: one-dimensional rotation curve of the galaxy with the best-fit rotational model (where appropriate) overlaid as contours. Right: one-dimensional distribution of linewidths. The solid bar in the left-hand panel denotes the amplitude of the seeing disc, converted to physical scale at the redshift of each galaxy.

Of course, the spatial distribution of an equivalent width depends on its intrinsic spatial distribution as well as the overall distribution of the continuum light (or galaxy surface brightness profile) after convolving with the seeing (e.g. Pracy et al. 2005). In particular, for a galaxy with a centrally concentrated equivalent width profile, the steeper the continuum surface brightness profile, the more the stellar population in the galaxy centre will contaminate the outer regions after convolution with the seeing (see Pracy et al. 2010, for a detailed discussion). In principle, this could give a correlation between $v/\sigma$ and the extent of the A-star light distribution. Whilst full modelling to recover the intrinsic equivalent width distribution is beyond the scope of this paper, we can perform a simple test of whether this effect is likely to significantly contribute our results. We therefore construct 10^4 mock GMOS data cubes with variable amplitude velocity fields, continuum emission and H$\delta$ equivalent width profiles (which broadly bracket the range of our observations) and convolve these with the seeing. We note that in these simulations, the extent of the continuum emission and H$\delta$ profiles are independent. We measure the ‘true’ $v/\sigma$ and ratio of $A(H\delta > 6)/A(gal)$ as well
as the same quantities from the recovered cubes after convolution with the seeing and surface brightness limits appropriate for our observations. Allowing the spatial extent of the A stars to vary from ‘unresolved’ to three times the continuum emission, we find that the change in slope of $\Delta \log(v/\sigma)/(\Delta A(H_\delta > 6\AA/\text{A(gal)})$ is $-0.04$ to $-0.2$ depending on choice of parameters. The gradient of the correlation seen in Fig. 5 is $\Delta \log(v/\sigma)/(\Delta A(H_\delta > 6\AA/\text{A(gal)}) = -2.2 \pm 0.3$. Thus the underlying A-star distribution, continuum surface brightness profile and seeing may contaminate the correlation between the extent of the A stars and dynamics seen in Fig. 5, but it seems unlikely that this effect alone can account for the correlation we see. We also note that, in general, rotation-dominated galaxies have shallower surface brightness profiles than pressure-supported galaxies (e.g. Courteau 1997). This means there is a physical correlation between the two which may lead to an indirect correlation between $v/\sigma$ and $A(H_\delta)/A(\text{gal})$. However, clearly, observations of a well-matched control sample of non-E+A galaxies around the $H_\delta$ are required to examine whether the correlation in Fig. 5 is a result of the processes which drive the E+A galaxies. Nevertheless, for our sample the ‘slow rotators’ tend to have the most widespread A-star populations, while the ‘fast rotators’ have the more compact A-star distributions.

4 DISCUSSION

There are at least three important unresolved problems regarding E+A galaxies. (1) From where in the galaxy does the unusual spectrum originate? (2) What star formation history leads to the spectral characteristics of these galaxies? (3) What triggers the recent change in star formation history? To address these issues, we have performed three-dimensional spectroscopy of a sample of 11 massive E+A galaxies selected from the SDSS for their unusually strong $H_\delta$ equivalent widths but weak [O II] emission, suggesting that star formation in these galaxies was recently truncated, possibly following a starburst.

Our sample was selected to span a range of morphology and environment and hence is diverse by design. Moreover, with a limited sample size of 11 galaxies it is not possible to answer any of these questions definitively, especially since there may be more than one way to form an E+A galaxy. Nevertheless, there are several important results that are generic within our sample. First, we note that we see no strong correlation of the galaxy dynamics with the K-band morphology, or the environment of the galaxy with their kinematics or morphology of the nebular emission or A stars; the two E+A galaxies in dense environments do not appear to stand out from the rest of the sample in any of their properties, so we discuss the properties of the ensemble below.

To investigate the likely age of the starbursts and hence star formation history, we estimate the age of the A-star populations by comparing the average galaxy spectrum for each galaxy with the spectral library of Jacoby, Hunter & Christian (1984). We use the strength of the Ca K $3933\,\AA$ absorption line (which is especially sensitive to age; Rose 1985), together with the equivalent width of the Ca H $3969\,\AA$ (and H e) absorption lines. For all of the E+A galaxies in our sample, the spectra best resemble that of an A5 $\pm 1$ star, with an effective temperature of $8160\,K$. Interpolating the theoretical isochrones from Bertelli et al. (1994), this suggests a luminosity-weighted age of $0.8 \pm 0.1\,\text{Gyr}$ [for the main-sequence (MS) turn-off at A8] although we caution that stellar masses and metallicity make this estimate uncertain. Nevertheless, this is also supported by the broad-band colours which are well matched by a model in which an instantaneous starburst (which accounts for 10 per cent by mass) is superposed upon the continuum above the 4000-$\AA$ break (Fig. 2).

Thus, the colours and line strengths we observe can best be reproduced with $\sim 10$ per cent (by mass) starbursts on top of an old population (e.g. Shioya, Bekki & Couch 2004; Balogh et al. 2005). It is also evident that any residual star formation in the galaxy is negligible: the median stellar mass of the E+A galaxies in our sample is $8 \pm 2 \times 10^{10}\,M_\odot$, suggesting an average burst mass of $M_{\text{burst}} \sim 1.0 \times 10^{10}\,M_\odot$. The median SFR of galaxies in our sample (including limits) is $\text{SFR}(\text{O} \text{[II]}) = 0.13 \pm 0.03\,M_\odot\,\text{yr}^{-1}$, less than 2 per cent of the past average SFR. Thus, the presence of even moderately strong $H_\delta$ absorption ($> 3\,\sigma$) or so indicates that the SFR at truncation was at least comparable to the past average rate. Integrating the current SFR over 1 Gyr we derive a mass of $\sim 1 \times 10^9\,M_\odot$, or only $\sim 0.1$ per cent of the total galaxy mass. Since the mass involved in the burst is at least $\sim 10$ per cent of the mass of the stellar remnant, this suggests that the main progenitors of our E+A galaxies had significantly higher SFRs than seen today, and were most likely blue-cloud, star-forming systems.

The A stars are widely distributed, with an average $A(H_\delta > 6\,\AA) = 4.0 \pm 2.0\,\text{kpc}^2$, and a range of $A(H_\delta > 6\,\AA) = 0.8 – 15.2\,\text{kpc}^2$ (see also Swinbank et al. 2005; Goto et al. 2008; Pracy et al. 2009). On average, the strongest $H_\delta$ $[W_c(H_\delta) > 6\,\AA]$ covers $33 \pm 5$ per cent of the area defined by the detected continuum above the 4000-$\AA$ break. It is also interesting that none of our galaxies shows regions with very strong absorption, $W_c(H_\delta) > 10\,\AA$, at least on $\sim$ kpc scales. Such strong absorption would be an unambiguous indication of a significant enhancement of the SFR in the past $< 1$ Gyr; weaker lines can be formed via truncation of star formation in normal, gas-rich galaxies (e.g. Balogh et al. 1999). Moreover, it is also clear that in approximately half of the sample, the A stars (which give rise to the absorption) and nebular emission are not co-located, suggesting that the newest stars are often forming in a different place.
than those that formed \( \lesssim 1 \) Gyr ago. Thus, together with the burst mass to stellar mass ratio, this suggests that there needs to have been a recent, dramatic change in star formation history, coordinated over kpc scales in these galaxies.

The galaxies display a range of kinematic properties: all 11 have velocity gradients in their stellar kinematics and we derive a median luminosity-weighted \( v\sin(i)/\sigma = 0.8 \pm 0.4 \) for the whole sample (with a range of 0.2–5.5). Fig. 7 shows that there is a weak correlation between absolute magnitude and \( v/\sigma \), such that the lower luminosity systems have an increasing rotational support. The median \( \lambda_R \) for the E+A galaxies in our sample is \( \lambda_R = 0.35 \pm 0.1 \), and all of the E+A galaxies studied here have \( \lambda_R > 0.1 \), which is consistent with that found for a comparably bright sample of elliptical galaxies and passive S0’s which are also well matched in their average stellar masses (Emsellem et al. 2007).

Although we have shown that the E+A phase is a galaxy-wide phenomenon, conclusively addressing what triggers the E+A phase is more difficult, especially since there are a number of processes which could be responsible at once (e.g. strong interactions or mergers, possibly proceeded by AGN activity). However, we can search for global trends in the A-star population with other galaxy properties in order to search for the major contributing factor. We find that dynamics are correlated with the spatial extent of the A stars such that the ‘slow rotators’ have the most widespread A-star populations, while the ’fast rotators’ have compact A-star distributions. Since the E+A galaxies in our sample do not preferentially lie in dense/cluster environments, it is unlikely that any truncation of star formation was due to ram pressure stripping of the gas.

It is interesting to note that theoretical models which aim to trace the formation and evolution of early-type galaxies have shown that there are two primary factors which determine the structure of high-z Early-type galaxies. More anisotropic and slowly rotating galaxies result from predominantly collisionless major mergers, while faster rotating galaxies are produced by more gas-rich mergers, where dissipation plays an important role (Kormendy & Djorgovski 1989; Bender, Burstein & Faber 1992; Faber et al. 1997). The second major factor which drives the galaxy structure is the mass fraction of the merger components. Unequal-mass galaxy mergers tend to produce more rapidly rotating galaxies than mergers with comparable mass ratios (Naab, Burkert & Hernquist 1999; Naab & Burkert 2003). Whichever process was responsible, the encounter has left residual rotational motion in at least 90 per cent of our sample, with A stars widely distributed within the ISM. Qualitatively, our results are consistent with models of unequal-mass, gas-rich mergers.

Finally, it is interesting to examine whether activation of a central black hole could be related to the end of star formation (Yan et al. 2006). As described in Section 2.2, four of our galaxies are detected in FIRST with 1.4-GHz fluxes in the range 0.5–2.2 mJy, corresponding to luminosities of \( L_{1.4} = 8–15 \times 10^{21} \) W Hz\(^{-1}\). All four galaxies with FIRST detections show [O III] and (weaker) [O II] emission. For these galaxies, the centroid of the nebular emission appears offset from the continuum by 0.4–1 kpc, but this is consistent with the rest of the sample. Thus, if there is an AGN in these galaxies, it does not appear to be the sole contributor to the nebular emission. Moreover, if the radio emission were due to star formation, then the instantaneous SFR would be \( 6–11 \) M\(_{\odot}\) yr\(^{-1}\) (Helou, Soifer & Rowan-Robinson 1985) a factor \( \sim 60 \) times that inferred from the nebular emission lines. This could be possible if the galaxies are exceptionally dusty, although this seems unlikely given the blue \((r - i)\) colours; we would expect dusty, star-forming galaxies to lie on the upper edge of the blue cloud in Fig. 2, with \((r - i) > 0.4\) (Wolf, Gray & Meisenheimer 2005; Balogh et al. 2009). Thus, it is most likely that the radio emission arises from an AGN. In all the key figures, we have highlighted the E+A galaxies in our sample that have \( L_{1.4} \gtrsim 10^{22} \) W Hz\(^{-1}\). In all cases, their properties appear to be consistent with those of the rest of the sample. Thus, although we only have small number statistics, our results suggest that AGN activity in E+A galaxies does not play a dramatic role in defining the spatial or kinematic properties of the A stars, and/or that its effects are short-lived.

5 CONCLUSIONS

We have investigated the spectrophotometric properties of a sample of 11 E+A galaxies in the redshift range 0.12 < z < 0.07 in order to provide constraints on the physical processes which cause the unusual post-starburst signatures. Although our sample selection was heterogeneous (by design we selected E+A galaxies with a range of morphologies and environments), there are several important results which can be summarized as follows.

(i) The strongest H\(\delta\) (\(W_\delta > 6 \) Å) tends to widely distributed within the galaxies, on average covering \( \sim 33 \) per cent of the galaxy image and extending over areas of \( 1–15 \) kpc\(^2\). This suggests that the characteristic E+A signature is a property of the galaxy as a whole and not due to a heterogeneous mixture of stellar populations.

(ii) In approximately half of the sample, the A stars, nebular emission and continuum are not co-located. The average offset between the H\(\delta\) and nebular emission is \( \Delta_{R_{H\delta}-gas} = 1.0 \pm 0.2 \) kpc, but can be as large as 1.7 kpc. This suggests that at least in some cases, the new OB stars are either displaced from the older A stars or that the extinction is not uniform. In either case, this helps to reconcile the fact that both emission and strong absorption can be seen in the same, integrated galaxy spectrum. The two populations do not generally arise from exactly the same place.

(iii) The colours and line strengths of these E+A galaxies are consistent with a \( \sim 10 \) per cent (by mass) starburst on top of an old population suggesting an average burst mass of \( M_{burst} \sim 10^{10} \) M\(_{\odot}\). The current (residual) star formation is unable to account for the burst mass. This indicates that the main progenitors of most of our E+A galaxies were blue-cloud, star-forming galaxies in which there was a recent, dramatic change in star formation history, coordinated over kpc scales.

(iv) The kinematics show that some level of rotation in the A-star population is common, and the dynamics (\( v/\sigma \) and \( \lambda_R \)) are consistent with those measured in comparably bright early types recently studied by Emsellem et al. (2007), which are also well matched in their average stellar masses. Whichever process was responsible for causing the starburst and subsequent truncation, the encounter has left residual rotational motion in at least 90 per cent of the sample, with A stars widely distributed within the ISM.

(v) We find that the A-star dynamics are correlated with their spatial extent such that the ‘slow rotators’ have the most widespread A-star populations, while the ‘fast rotators’ have compact A-star distributions.

(vi) We also find that the fraction of galaxies which have radio emission suggestive of low luminosity AGN is 20–40 per cent, a factor \( \sim 8 \) times higher than expected given their stellar masses, indicating a high AGN fraction of E+A galaxies. However, although our sample is limited to only 11 galaxies, the kinematics and spatial
distribution of the stars in the radio-detected subsample do not stand out from the radio-undetected systems, suggesting that AGN feedback in E+A galaxies does not play a dramatic role in defining their properties, and/or that its effects are short.

Whichever process was responsible for the E+A signatures, rotational motion is seen in the majority of our sample, and their dynamics, stellar masses and luminosities are consistent with those expected for the progenitor population of (at least a subset of) representative, early-type galaxies. Overall, these observations provide new and detailed constraints on the kinematics and spatial distribution of the A stars and gas in E+A galaxies, and hence constraints on galaxy formation models which aim to test the variety of physical mechanisms which trace the evolutionary sequence linking gas-rich, star-forming galaxies to quiescent spheroids.

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